

# Nonthermal and thermal emission from the supernova remnant RX J1713.7-3946

E.G. Berezhko<sup>1</sup> and H.J.Völk<sup>2</sup>

<sup>1</sup> Yu.G. Shafer Institute of Cosmophysical Research and Aeronomy, 31 Lenin Ave., 677980 Yakutsk, Russia  
e-mail: berezhko@ikfia.ysn.ru

<sup>2</sup> Max Planck Institut für Kernphysik, Postfach 103980, D-69029 Heidelberg, Germany  
e-mail: Heinrich.Voelk@mpi-hd.mpg.de

Received month day, year; accepted month day, year

## ABSTRACT

**Aims.** A nonlinear kinetic theory of cosmic ray (CR) acceleration in supernova remnants (SNRs) is employed to investigate the properties of SNR RX J1713.7-3946.

**Methods.** Observations of the nonthermal radio and X-ray emission spectra as well as the H.E.S.S. measurements of the very high energy  $\gamma$ -ray emission are used to constrain the astronomical and CR acceleration parameters of the system. It is argued that RX J1713.7-3946 is a core collapse supernova (SN) of type II/Ib with a massive progenitor, has an age of  $\approx 1600$  yr and is at a distance of  $\approx 1$  kpc. It is in addition assumed that the CR injection/acceleration takes place uniformly across the shock surface for this kind of core collapse SNR.

**Results.** The theory gives a consistent description for all the existing observational data, including the nondetection of thermal X-rays and the spatial correlation of the X-ray and  $\gamma$ -ray emission in the remnant. Specifically it is shown that an efficient production of nuclear CRs, leading to strong shock modification and a large downstream magnetic field strength  $B_d \approx 140 \mu\text{G}$  can reproduce in detail the observed synchrotron emission from radio to X-ray frequencies together with the  $\gamma$ -ray spectral characteristics as observed by the H.E.S.S. telescopes.

**Conclusions.** The calculations are consistent with RX J1713.7-3946 being an efficient source of nuclear cosmic rays.

**Key words.** ISM: cosmic rays – acceleration of particles – shock waves – stars: supernovae: individual:SNR RX J1713.7-3946 – radiation: emission mechanisms: nonthermal– gamma-rays: theory

## 1. Introduction

RX J1713-3946 is a shell-type supernova remnant (SNR) located in the Galactic plane that was discovered in X-rays with *ROSAT* (Pfeffermann & Aschenbach, 1996). Subsequent studies of this SNR with the *ASCA* satellite by Koyama et al. (1997) and Slane et al. (1999), and later with *XMM* by Cassam-Chenai et al. (2004), have not positively identified any thermal X-ray emission and therefore tentatively concluded that the observable X-ray emission is entirely non-thermal.

The radio emission on the other hand is quite weak. Only part of the remnant shell could be detected in radio synchrotron emission up to now, with a poorly known spectral form (Lazendic et al., 2004). For the spatially integrated radio flux Aharonian et al. (2006) estimated about twice the value found by Lazendic et al. (2004). Most recently Acero et al. (2009) derived a total flux value of  $22\text{Jy} < S < 26\text{Jy}$  at 1.4 GHz, still about two times higher than estimated by Aharonian et al. (2006). As far as the estimate of the magnetic field strength is concerned, this higher estimate of the radio

synchrotron flux increases the magnetic field estimate (see below).

RX J1713-3946 was also detected in very high energy (VHE:  $> 100$  GeV)  $\gamma$ -rays with the *CANGAROO* (Muraishi et al., 2000; Enomoto et al., 2002) and *H.E.S.S.* telescopes (Aharonian et al., 2004, 2006, 2007). Especially the latter, very detailed observations show a clear shell structure at TeV energies which correlates well with the *ASCA* contours.

In a first theoretical paper on this source (Berezhko & Völk, 2006) we have investigated the acceleration of CR electrons and protons in detail, using nonlinear kinetic theory (Berezhko et al., 1996; Berezhko & Völk, 2000). Observations of the nonthermal radio and X-ray emission spectra as well as the H.E.S.S. measurements of the very high energy  $\gamma$ -ray emission were used to constrain the astronomical and the particle acceleration parameters of the system. Under the assumption that RX J1713.7-3946 is the remnant of a core collapse supernova (SN) of type II/Ib with a massive progenitor, the theory gives indeed a consistent description for the existing data on the nonthermal emission. Specifically the VHE data from H.E.S.S. were shown to be best explained as hadronic  $\gamma$ -ray emission, where the magnetic field amplification strongly

depresses the inverse Compton and Bremsstrahlung fluxes. Subsequently (Berezhko & Völk, 2008), we have analyzed the spatial correlation between the synchrotron and the VHE emission that has been used before as an argument for a leptonic origin of the VHE emission (Katz & Waxman, 2008; Plaga, 2008). It was argued that correlated density and magnetic field strength variations lead rather naturally to a spatial correlation of the hadronic VHE emission with the synchrotron emission. Similar arguments were brought forward by Tanaka et al. (2008). Also the recent broad-band X-ray synchrotron measurements (Uchiyama et al., 2007; Takahashi et al., 2008) with the Suzaku instrument were compared to the theoretical synchrotron spectra and found to be quite consistent with a hadronic model. A purely leptonic model on the other hand (e.g. Porter et al., 2006), where magnetic field amplification is not expected, was shown to be inconsistent with these observations. (Acero et al. (2009) have also considered a leptonic origin of the  $\gamma$ -ray emission as a viable possibility. This has to be compared with the results of the present paper.)

These theoretical considerations assumed that only over part of the shock surface injection of suprathermal nuclear ions occurs effectively, because over other parts the magnetic field is essentially parallel to the shock surface. In fairly regular SNRs like those of Type Ia explosions into a presumably uniform circumstellar environment, this is an important effect, as for instance the case of SN 1006 shows. For a blast wave propagating into the highly turbulent shell of a stellar wind bubble from a massive progenitor star on the other hand, the magnetic flux tubes in the shell are presumably so slender and irregular that energetic particles can cross effusively into regions where suprathermal injection is depressed, yet acceleration is possible. As a result shock modification by the pressure gradient of the accelerating particles becomes possible everywhere on the shock surface. This reasoning follows the discussion in Berezhko et al. (2009) with the conclusion that SNRs propagating into wind bubbles should have their forward shock modified everywhere. Technically this implies that the SNR becomes spherically symmetric and the correction factor  $f_{\text{re}} < 1$  (Völk et al., 2003) on the spherically symmetric solution reaches unity. The ubiquitous shock modification has also implications for the thermal emission, because the thermal gas is then heated only in the subshock and the average gas density required for a given hadronic  $\gamma$ -ray emission is reduced.

Contrary to the previous studies (Berezhko & Völk, 2006, 2008) of RX J1713.7-3946 we therefore adopt here the approximation of spherically symmetric CR injection/acceleration with  $f_{\text{re}} \approx 1$  (see also Zirakashvili, 2009). It is demonstrated that also in this approximation a consistent solution can also be found. In addition a rough estimate of the thermal X-ray emission from this object shall be given. Even though the error in this estimate might be quite large, the nominal result is consistent with the present non-detection of thermal X-rays (e.g. Acero et al., 2009). Finally, the observed correlation of the X-ray and the VHE  $\gamma$ -ray emission is discussed in a generalized form.

## 2. Results

As in the previous studies (Berezhko & Völk, 2006, 2008) a source distance of  $d = 1$  kpc is adopted here. For the present angular source size of about 60 arcmin this implies a SNR radius  $R_s \approx 10$  pc.

For the SNR age the value  $t_{\text{sn}} = 1612$  yr is used, consistent with the hypothesis of Wang et al. (1997) that RX J1713-3946 is the remnant of the AD393 “guest star”.

It had been argued (Berezhko & Völk, 2006) that the existing data are only consistent with RX J1713.7-3946 being a core collapse SN of type II/Ib. Those progenitor stars of core collapse SNe that emit intense winds are massive main-sequence stars with initial masses  $M_i \gtrsim 15 M_\odot$ . During their evolution in a surrounding uniform ISM of gas number density  $\varrho_0 = m_p N_{\text{ISM}}$ , the wind termination shock creates a hot bubble of shocked wind material, surrounded by a turbulent shell of shocked ISM behind an outer forward shock that communicates the internal overpressure to the environment.

The gas number density distribution  $N_g(r) = \varrho(r)/m_p$  is assumed to have the form

$$N_g = 0.008 + 0.24[r/(10 \text{ pc})]^{12} \text{ cm}^{-3} \quad (1)$$

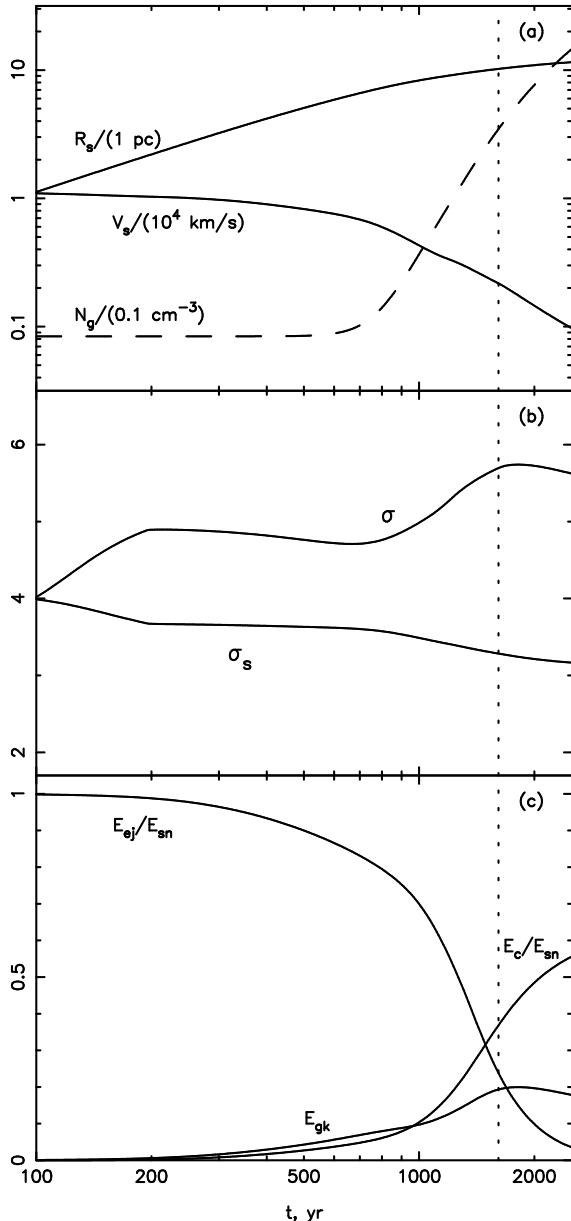
This density profile corresponds to a gas number density  $N_g = 0.25 \text{ cm}^{-3}$  at the current SN shock position, which is a factor  $\sim 5$  lower than in Berezhko & Völk (2006).

Under these conditions a SN explosion energy  $E_{\text{sn}} = 1.3 \times 10^{51} \text{ erg}$  and an ejecta mass  $M_{\text{ej}} = 3 M_\odot$  lead to a good fit for the observed SNR properties. To determine the explosion energy directly from the data the value of the shock speed  $V_s$  would have been needed. However, this quantity is not measured. The ejecta energy  $E_{\text{ej}}$  is still about  $0.25 E_{\text{sn}}$ . Together with the kinetic energy of the shocked gas  $E_{\text{gk}} \approx 0.2 E_{\text{sn}}$  this makes up about half of the total energy  $E_{\text{sn}}$  (see Fig.1).

The spectral fit, also for the  $\gamma$ -ray data, can be achieved with the values  $\eta = 5 \times 10^{-4}$  of the proton injection rate and  $B_d = 142 \mu\text{G}$  for the downstream magnetic field strength. Such values of  $B_d$  are significantly higher than typical ambient magnetic fields, also for the considered bubble wall in the dense environment (the gas density at the shock is still a factor  $\sim 100$  lower than in that environment which implies a strongly decompressed field upstream of the SNR shock compared to the value in the ambient medium). The high  $B_d$ -value must be attributed to field amplification at the shock front due to the strong wave production by the acceleration of CRs far into the nonlinear regime (Lucek & Bell, 2000; Bell, 2004)<sup>1</sup>

It is assumed that also electrons are injected into the acceleration process at the shock front. Since the details of the electron injection process are poorly known, the electron injection rate is chosen such that the electron to proton ratio  $K_{\text{ep}}$  (which is defined as the ratio of their distribution functions at all rigidities where the protons are already relativistic and the electrons have not been yet cooled radiatively) is a constant to be determined from the synchrotron observations. Clearly,

<sup>1</sup> For direct observational evidence regarding magnetic field amplification in RX J1713.7-3946 see Völk et al. (2005); Ballet (2006); Berezhko & Völk (2006).



**Fig. 1.** (a) Shock radius  $R_s$ , shock speed  $V_s$ , and preshock gas density  $N_g$ ; (b) total shock ( $\sigma$ ) and subshock ( $\sigma_s$ ) compression ratios; (c) ejecta ( $E_{ej}$ ) gas kinetic ( $E_{gk}$ ) and CR ( $E_c$ ) energies as a function of time, normalised to the total mechanical SNR energy  $E_{sn}$ . The vertical dotted line marks the current evolutionary epoch.

from the point of view of injection/acceleration theory,  $K_{ep}$ , together with  $B_0 \equiv B_d/\sigma$  and  $\eta$ , must be treated as theoretically not very well constrained parameters to be *quantitatively* determined by comparison with the available synchrotron and  $\gamma$ -ray data. Here  $\sigma$  is the overall shock compression ratio.

In the present case the following parameter values are obtained by iteration:  $\eta = 5 \times 10^{-4}$ ,  $B_0 = 25 \mu\text{G}$ ,  $K_{ep} = 1.3 \times 10^{-3}$ . The gas dynamic variables at the present epoch are  $\sigma = 5.7$ ,  $\sigma_s = 3.3$ , and  $V_s \approx 2200 \text{ km sec}^{-1}$ , where  $\sigma_s$  and  $V_s$  denote the subshock compression ratio and the overall shock velocity, respectively (Fig.1).

The corresponding solutions of the dynamic equations at each instant of time yield the CR spectrum and the spatial distributions of CRs and thermal gas, and therefore also the expected fluxes of nonthermal emission produced by the accelerated CRs.

The overall broadband spectral energy distribution (SED), expected at the current evolutionary epoch, is displayed in Fig.2, together with the experimental data from ATCA at radio wavelengths, as estimated for the full remnant by Acero et al. (2009), the X-ray data from ASCA (Aharonian et al., 2006) and Suzaku (Uchiyama et al., 2007), and the H.E.S.S.  $\gamma$ -ray data (Aharonian et al., 2007).

The comparison with our earlier theoretical spectrum (Berezhko & Völk, 2008) shows that also a model with full spherical symmetry can successfully describe the spatially integrated nonthermal emission properties of this object.

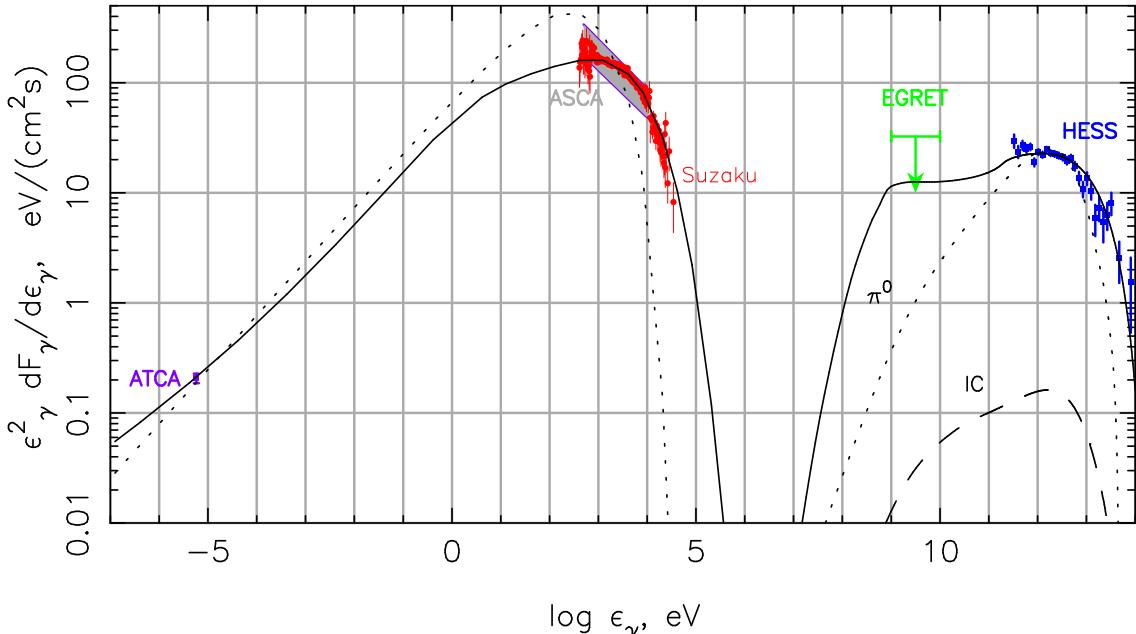
At energies  $\epsilon_\gamma < 1 \text{ TeV}$  the theoretical  $\gamma$ -ray spectrum is as hard as  $dF_\gamma/d\epsilon_\gamma \propto \epsilon_\gamma^{-1.8}$ , whereas for  $\epsilon_\gamma > 10 \text{ TeV}$  it has a smooth cutoff. It should be noted that the  $\gamma$ -ray cutoff energy  $\epsilon_\gamma^{\max} \approx 0.1 c p_{\max}$  is sensitive to the magnetic field strength  $B_d$ , since the proton cutoff momentum has a dependence  $p_{\max} \propto R_s V_s B_d$  (Berezhko, 1996).

The theoretical results presented here are fully consistent with a dominantly hadronic origin of the observed TeV  $\gamma$ -rays. At the present epoch the SNR has already converted  $\approx 35\%$  of  $E_{sn}$  into accelerated nuclei.

## 2.1. Estimate of the thermal X-ray emission

Like for the case of RX J0852.0-4622 (Berezhko et al., 2009) in addition the expected flux of thermal X-rays is estimated here. This is only possible in a very approximate way, even if we shall disregard the ejecta emission with the argument that the ejected mass in the present phase is small compared to the swept-up mass. The reason is the explosion into the wind bubble, creating a rather different thermodynamical structure from that of a classical Sedov solution for a SN explosion into a uniform medium, exclusively until now considered in the literature. Yet the remnant of RX J1713.7-3946 is well past the sweep-up phase and has entered a quasi-Sedov phase in the stellar wind shell, i.e. a roughly self-similar evolutionary phase, modified by strong particle acceleration relative to a purely gas dynamic evolution. Other estimates of the thermal emission from RX J1713.7-3946 have been made by Katz & Waxman (2008) and Drury et al. (2008), effectively assuming an explosion into a uniform medium.

The approximation adopted here is the following. The present bubble case is compared with a SNR in a uniform medium in the classical Sedov phase without any CR acceleration, making four assumptions (i) the total hydrodynamic explosion energy is the same in both cases (ii) the shock velocity is the same (iii) the present gas density upstream of the shock is the same, and (iv) the two objects are at the same distance of 1 kpc. Then the results of Hamilton et al. (1983) for the thermal X-ray flux from a classical Sedov SNR are used, employing the emission measure of the bubble remnant instead of that of the classical Sedov remnant with the same four parameters above.



**Fig. 2.** Spatially integrated, overall nonthermal spectral energy distribution of RX J1713.7-3946. The *solid* curve at energies above  $10^7$  eV corresponds to  $\pi^0$ -decay  $\gamma$ -ray emission, whereas the *dashed* curve indicates the Inverse Compton (IC) emission. The dotted line corresponds to the test particle limit which implies insignificant proton acceleration and magnetic field amplification (see Berezhko & Völk, 2008, for the details). The ATCA radio data, as derived by Acero et al. (2009), the ASCA X-ray data (cf. Aharonian et al., 2006), the Suzaku X-ray data (Uchiyama et al., 2007), and the 2006 HESS  $\gamma$ -ray data (Aharonian et al., 2007) are also shown. The EGRET upper limit for the RX J1713.7-3946 position (Aharonian et al., 2006) is included as well.

This means that the X-ray emissivity of the remnant is reduced by the ratio  $R_{\text{em}} = EM_b/EM_s$  of the emission measure  $EM_b$  for our bubble solution to the emission measure for the classical Sedov solution  $EM_s$  which corresponds to a uniform ambient gas density  $N_g(R_s)$  (Berezhko et al., 2009).

The relation

$$\frac{T_{\text{sub}}}{T_s} = \frac{[(\sigma_s - 1)(\gamma + 1) + 2](\gamma + 1)^2}{4\gamma(\gamma - 1)\sigma^2} \quad (2)$$

is used, with  $\gamma = 5/3$ , to determine the subshock temperature  $T_{\text{sub}}$ , relevant for the thermal emission flux for the modified shock.  $T_s = 10^7 (V_s/(839 \text{ km/s}))^2 \text{ K} \approx 6.9 \times 10^7 \text{ K}$  denotes the gas temperature of an unmodified shock with  $V_s \approx 2200 \text{ km/s}$ .

Using the differential thermal X-ray model spectra  $dF/d\epsilon \approx 3 \times 10^{-5}$  photons/(keV cm<sup>2</sup>s) from Hamilton et al. (1983) (for  $T_e \neq T_i$ , see their Fig.2) for their  $\eta = N_{\text{H}}^2 E_{\text{sn}} = 10^{49} \text{ erg cm}^{-6}$  with account of the scaling  $dF/d\epsilon \propto \eta$  with our model numbers  $\eta \approx 8 \times 10^{49} \text{ erg cm}^{-6}$ , multiplying it by the factor  $\theta^2 [E_{\text{sn}}/(10^{51} \text{ erg})]^{-1/2}$ , as required, where  $\theta = 54.5 \text{ arcmin}$  for the angular size of the classical Sedov remnant corresponding to RX J1713.7-3946 and multiplying it also by the factor  $R_{\text{em}} = 0.46$ , results in a thermal spectral energy density  $\epsilon^2 dF/d\epsilon \approx 288 \text{ eV cm}^{-2} \text{ sec}^{-1}$  for  $\epsilon = 1 \text{ keV}$ .

This result has not yet taken into account that the actual postshock gas temperature  $T_{\text{sub}}$  is by a factor  $\approx 0.41$  lower than the temperature  $T_s$  used in the above estimate. As discussed in Berezhko et al. (2009) this roughly corresponds to a reduction of the parameter value  $\eta \approx 3.3 \times 10^{49} \text{ erg cm}^{-6}$  and thus to a reduction of the thermal flux at 1 keV to  $\epsilon^2 dF/d\epsilon \approx 118 \text{ eV cm}^{-2} \text{ sec}^{-1}$ .

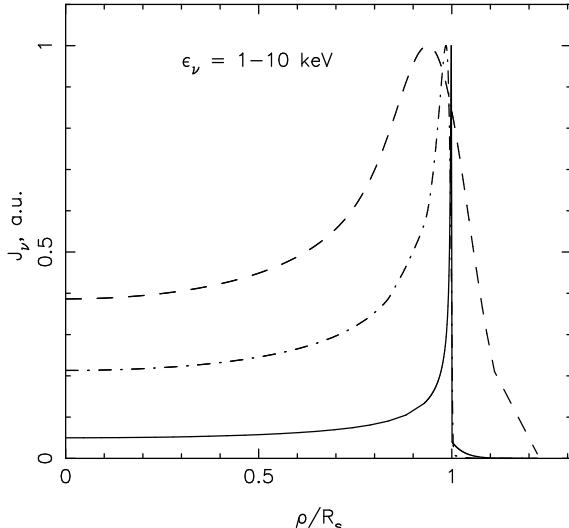
These numbers must be compared with the observed non-thermal X-ray energy flux at 1 keV (see Fig.1)  $\epsilon^2 dF/d\epsilon \approx 200 \text{ eV cm}^{-2} \text{ sec}^{-1}$ . Therefore, at 1 keV, the thermal energy flux comes out to be smaller than the nonthermal flux by a factor of about 2. This does not contradict the non-detection of thermal X-rays from RX J1713.7-3946.

## 2.2. Spatial correlation of the nonthermal X-ray fluxes with the $\gamma$ -ray fluxes

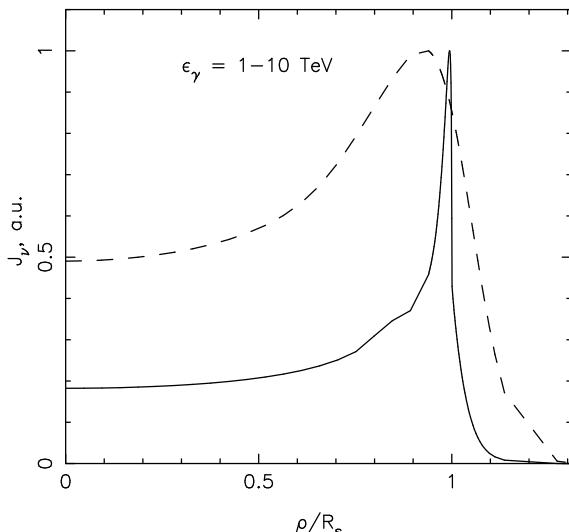
Note that Acero et al. (2009), who find a positive spatial non-linear correlation  $F_X \propto F_\gamma^\alpha$  with  $\alpha = 2.41$  between the flux  $F_\gamma$  of  $\gamma$ -rays with energies  $\epsilon_\gamma = 1 - 10 \text{ TeV}$  and the flux of non-thermal X-rays  $F_X$  with energies  $\epsilon_\gamma = 1 - 10 \text{ keV}$ , concluded that such a correlation is in favor of a leptonic nature of the TeV-emission. Such a positive correlation is naturally expected within the present approach.

### 2.2.1. Correlation of the radial profiles

In terms of the spherically symmetric model, discussed here, the correlation between X-ray and  $\gamma$ -ray fluxes means that the projected radial profiles  $J_X(\rho)$  and  $J_\gamma(\rho)$  of these emissions as functions of projected radial distance  $\rho$  have a similar shape. At first sight this is a contradiction to this model: as can be seen from Figs.3 and 4, the calculated profiles are significantly different. Both profiles have a sharp peak just behind the shock front, but the X-ray profile is by a factor of about 3 thinner than the  $\gamma$ -ray profile. However, when smoothed with the point



**Fig. 3.** The X-ray emissivities for the energy  $\epsilon_\gamma = 1 - 10$  keV as a function of projected, normalized radial distance  $\rho/R_s$ . The calculated radial profile is represented by the *solid line*. The *dashed line* represents the calculated profile convolved with the point spread function of Gaussian width  $0.05^\circ$ .



**Fig. 4.** The  $\gamma$ -ray emissivities for the  $\gamma$ -ray energy  $\epsilon_\gamma = 1 - 10$  TeV as a function of projected, normalized radial distance  $\rho/R_s$ . The calculated radial profile is represented by the *solid line*. The *dashed line* represents the calculated profile convolved with the point spread function of Gaussian width  $0.05^\circ$ .

spread function of width  $\sigma_0 = 0.05^\circ$ , corresponding to the H.E.S.S. angular resolution, these profiles become very similar. (Acero et al., 2009) compared XMM measurements of the X-ray emission smoothed with a point spread function of width  $\sigma = 0.05^\circ$ . It can therefore be concluded that the similarity of the X-ray and the  $\gamma$ -ray radial profiles is a trivial consequence of the smoothing of two differently sharp peaks by a point spread function whose width is considerably larger than the widths of either one of the emission peaks.

The experimentally observed correlation according to Fig.9 of Acero et al. (2009) is rather close to the relation  $F_X = F_\gamma$

for all the brighter SNR regions. Some deviation from this relation is observed for six of the outer and dimmer regions, for which  $F_X < F_\gamma$ . This peculiarity can be also explained within the present model. The reason is that protons with an energy of about 100 TeV occupy a region, which is noticeably larger than the shock size  $R_s$ , due to their high diffusive mobility. The 10 TeV  $\gamma$ -ray emission which these protons produce has an even larger radial extension since the gas density strongly increases with radial distance. Such an effect is absent for the electrons of the same energy which produce the nonthermal keV-emission: as a result of their strong energy losses they occupy a region thinner than  $10^{-2}R_s$  around the shock front. Therefore, as can be clearly seen from Figs.3 and 4, the ratio  $F_\gamma/F_X$  is expected to increase as a function of radial distance for  $\rho/R_s > 1$ , as the data of Acero et al. (2009) are interpreted here.

## 2.2.2. Correlation of azimuthal variations

The measured X-ray and  $\gamma$ -ray emissions also undergo rather closely correlated azimuthal variations across the remnant. Such an effect can not be described within the present model model, because it is spherically symmetric. One can nevertheless attempt to interpret the correlated behaviour of the fluxes  $F_X$  and  $F_\gamma$  in different azimuthal sectors of the remnant. It is natural to assume that this variation of the remnant properties is due to variations of the ambient gas density  $\varrho$ . According to the model proposed here the SNR's present evolutionary phase is intermediate between the sweep-up and a quasi-Sedov phase. Since the ejecta kinetic energy together with the kinetic energy of the swept-up gas contain about half of the explosion energy, the SN shock can be treated approximately as a piston-driven shock. The speed of such a shock depends only weakly upon the upstream gas density. Therefore  $V_s$  can be approximated by a constant across the whole remnant.

The local hadronic  $\gamma$ -ray emission then varies like  $F_\gamma \propto \varrho E_c \propto \varrho$ , where the local CR energy content  $E_c$  in each angular sector is assumed to be the same everywhere, because the overall CR energy content  $E_c \approx 0.4E_{SN}$  has already reached the saturation level and therefore can not undergo significant variations. In addition, the highest energy CRs in the SNR have the highest diffusive mobility so that they tend to be distributed roughly uniformly across the remnant.

As it is clear from Fig.2, the synchrotron spectrum at high energies  $\epsilon_\gamma > 10$  eV is already considerably influenced determined by synchrotron losses of the emitting electrons. Since the energy of these electrons is rather rapidly and completely transformed into synchrotron emission, the flux  $F_X \propto K_{ep} \varrho V_s^3$  is determined by the available energy flux  $\rho V_s^3/2$  into the shock, an essential part of which is transformed at the shock front into the energy flux of high energy CRs. High energy electrons therefore accumulate a fraction  $K_{ep}$  of this flux. This flux hardly depends on the value of the interior magnetic field in all those SNR regions where the field is sufficiently high,  $B_d \gtrsim 140 \mu\text{G}$ . In this case the relation  $F_X \propto F_\gamma$  is expected for the brightest part of remnant. This is indeed observed (Acero et al., 2009). Since the energy density of the

amplified magnetic field  $B_d^2$  is phenomenologically known to be proportional to  $\varrho V_s^\beta$ , with  $2 \leq \beta \leq 3$  (Völk et al., 2005), it is expected to be proportional to the local upstream gas density  $\varrho$ . The magnetic field strength is therefore lower than the average value  $\approx 140 \mu\text{G}$  within those parts of the remnant, where the gas density is lower than the average value. Synchrotron losses become much smaller in these regions and therefore the expected synchrotron X-ray flux is approximately  $F_X \propto \varrho K_{\text{ep}} B_d^{3/2} \propto \varrho^{7/4}$ . For the dim part of the remnant, in which X-ray and  $\gamma$ -ray emission is considerably lower than on average, one then obtains  $F_X \propto F_\gamma^{7/4}$ . This is approximately consistent with the observational result.

### 3. Summary

The assumption that CR injection/acceleration takes place uniformly across the whole SN shock surface is consistent with the existing data.

The swept-up mass is so small in this case that, in a rough approximation, the estimated flux of thermal X-rays at 1 keV is lower than the nonthermal flux, consistent with the nondetection of thermal X-ray emission until now.

It is concluded that the present observational knowledge of SNR RX J1713.7-3946 can be interpreted by a source which ultimately converts more than 35% of the mechanical explosion energy into nuclear CRs. Also the observed high energy  $\gamma$ -ray emission of SNR RX J1713.7-3946 turns out to be primarily of hadronic origin.

*Acknowledgements.* The authors thank V.S. Ptuskin and V.N. Zirakashvili for discussions on the thermal emission properties. This work has been supported in part by the Russian Foundation for Basic Research (grants 06-02-96008, 07-02-0221). EGB acknowledges the hospitality of the Max-Planck-Institut für Kernphysik, where part of this work was carried out.

### References

Abbott, D.C. 1982, ApJ, 263, 723

Acero, F., Ballet, J., Decourchelle, A., et al. 2009, to be published in A&A; arXiv:0906.1073v1 [astro-ph.HE]

Aharonian, F. A., Akhperjanian, A., Aye, K.-M., et al. (H.E.S.S. Collaboration) 2004, Nature, 432, 75

Aharonian, F. A., Akhperjanian, A., Bazer-Bachi, A.R., et al. (H.E.S.S. Collaboration) 2006, A&A, 449, 223

Aharonian, F. A., Akhperjanian, A., Bazer-Bachi, A.R., et al. (H.E.S.S. Collaboration) 2007a, A&A, 464, 235

Ballet, J. 2006, Adv. Space Res., 37, 1902

Bell, A. R. 2004, MNRAS, 353, 550

Berezhko, E.G. 1996, Astropart. Phys., 5, 367

Berezhko, E.G., Elshin, V.K. & Ksenofontov, L.T. 1996 JETPh, 82, 1

Berezhko, E.G. & Völk, H.J. 2000, A&A, 357, 183

Berezhko, E. G. 2005, Adv. Space Res., 35, 1031

Berezhko, E.G. 2008, Adv. Space Res., 41, 429

Berezhko, E.G. & Völk, H.J. 2006, A&A, 451, 981

Berezhko, E.G. & Völk, H.J. 2008, A&A, 492, 695

Berezhko, E.G., Ksenofontov, L.T., & Völk, H.J. 2009, A&A, in press; arXiv:0906.3944v1 [astro-ph.HE]

Berezhko, E. G., Pühlhofer, G., & Völk, H. J. 2009, A&A in press; arXiv:0906.5158v1 [astro-ph.HE]

Cassam-Chenaï, G., Decourchelle, A., Ballet, J., et al. 2004, A&A, 427, 199

Drury, L.O'C., Aharonian, F.A., Malyshev, D. & Gabici, S. 2008, A&A, 496, 1

Enomoto, R., Tanimori, T., Naito, T., et al. 2002, Nature, 416, 823

Enomoto, R., Watanabe, S., Tanimori, T. et al. 2006, ApJ, 652, 1268

Fukui, Y., Moriguchi, Y., Tamura, K. et al. 2003, PASJ, 55, L61

Hamilton, A.J.S., Sarazin, C.L. & Chevalier, R.A. 1983, ApJS, 41, 115

Hiraga, J.S., Uchiyama, Y., Takahashi, T., et al. 2005, A&A, 431, 953

Katz, B. & Waxman, E. 2008, JCAP, 01, 1

Koyama, K., Kinagasa K., Matsuzaki, K., et al. 1997, PASJ, 49, L7

Lazendic, J. S., Slane, P.O., Gaensler, B.M., et al. 2004, ApJ, 602, 271

Lucek, S.G. & Bell, A.R. 2000, MNRAS, 314, 65

Moriguchi, Y., Tamura, T., Tawara, Y., et al. 2005, ApJ, 641, 947

Muraishi, H., Tanimori, T., & Yanagita, S. 2000, A&A, 354, L57

Pfeffermann, E. & Aschenbach, B. 1996, in: Röntgenstrahlung from the Universe. (Eds. H.U. Zimmermann, J. Trümper & H. Yorke (MPE Rep. 263, Garching), 267

Plaga, R. 2008, New Astronomy, 13, 73

Porter, T.A., Moskalenko, I.V., Strong, A.W. 2006, ApJ, 648, L29

Slane, P., Gaensler, B. M., Dame, T. et al. 1999, ApJ, 357, SL99

Takahashi, T., Tanaka, T., Uchiyama, Y., et al. 2008, PASJ, 60, S131

Tanaka, T., Uchiyama, Y., Aharonian, F.A., et al. 2008, to appear in ApJ, [arXiv:0806.1490 (astro-ph)]

Uchiyama, Y., Aharonian, F.A., Tanaka, T., et al. 2007, Nature, 449, 576

Völk, H.J., Berezhko, E.G. & Ksenofontov, L.T. 2003, A&A, 409, 563

Völk, H.J., Berezhko, E.G. & Ksenofontov, L.T. 2005, A&A, 433, 229

Wang, Z.R., Qu, Q.-Y., & Chen, Y. 1997, A&A, 318, L59

Zirakashvili, V.N. 2009, in "High Energy Gamma-Ray Astronomy" (Eds. F.A. Aharonian, W. Hofmann, F.M. Rieger), Melville, New York, 2009, AIP Conf. Proc. 1085, p. 129 ff.